

ON THE RELATIVISTIC PRECESSION AND OSCILLATION FREQUENCIES OF TEST PARTICLES AROUND RAPIDLY ROTATING COMPACT STARS

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ABSTRACT

Whether analytic exact vacuum(electrovacuum) solutions of the Einstein(Einstein-Maxwell) field equations can accurately describe or not the exterior spacetime of compact stars remains still an interesting open question in Relativistic Astrophysics. As an attempt to establish their level of accuracy, the radii of the Innermost Stable Circular Orbits (ISCOs) of test particles given by analytic exterior spacetime geometries have been compared with the ones given by numerical solutions for neutron stars (NSs) obeying a realistic equation of state (EoS). It has been so shown that the six-parametric solution of Pachón, Rueda, and Sanabria (2006) (hereafter PRS) is more accurate to describe the NS ISCO radii than other analytic models. We propose here an additional test of accuracy for analytic exterior geometries based on the comparison of orbital frequencies of neutral test particles. We compute the Keplerian, frame-dragging, as well as the precession and oscillation frequencies of the radial and vertical motions of neutral test particles for the Kerr and PRS geometries; then we compare them with the numerical values obtained by Morsink and Stella (1999) for realistic NSs. We identify the role of high-order multipole moments such as the mass quadrupole and current octupole in the determination of the orbital frequencies especially in the rapid rotation regime. The results of this work are relevant to cast a separatrix between black hole (BH) and NS signatures as well as probe the nuclear matter EoS and NS parameters from the Quasi-Periodic Oscillations (QPOs) observed in Low Mass X-Ray Binaries.

Subject headings: Exact solutions Einstein-Maxwell equations – Relativistic precession frequencies

1. INTRODUCTION

One of the greatest challenges of the general theory of relativity has been the construction of solutions to the Einstein-Maxwell field equations representing the gravitational field of compact stars such as neutron stars (NSs). Stationary axially symmetric spacetimes satisfy basic properties one expects for rotating objects, namely time symmetry and reflection symmetry with respect to the rotation axis (see e.g. Pachón and Sanabria-Gómez 2006). The simplest stationary axially symmetric exact exterior vacuum solution describing a rotating configuration is the well-known Kerr metric (Kerr 1963). The Kerr metric is fully described by two free parameters: the mass M and the angular momentum J of the object. However, it is known from numerical models that the quadrupole moment of rotating NSs deviates considerably from the one given by the Kerr solution $Q_{\text{Kerr}} = -J^2/(Mc^2)$ (see e.g. Laarakkers and Poisson 1999, for details).

In the mean time, a considerable number of analytic exterior solutions with a more complex multipolar structure than the one of the Kerr solution have been developed (see e.g. Manko *et al.* 1995, 2000; Stephani *et al.* 2003). Whether analytic exterior solutions are accurate or not to describe the gravitational field of compact stars is an interesting and

very active topic of research (see e.g. Stute and Camenzind 2002; Berti and Stergioulas 2004; Pachón *et al.* 2006, and references therein).

The accuracy of analytic solutions to describe the exterior geometry of a realistic rotating compact star has been tested by comparing physical properties, e.g. the radius of the Innermost Stable Circular Orbit (ISCO) on the equatorial plane and the gravitational redshift (see Sibgatullin and Sunyaev 1998; Berti and Stergioulas 2004; Pachón *et al.* 2006, for details). In order to do such a comparison, the free parameters (i.e. the lowest multipole moments) of the analytic exterior spacetime, are fixed to the corresponding lowest multipole moments given by numerical interior solutions of the Einstein equations, for NS realistic models (see e.g. Berti and Stergioulas 2004).

Following such a procedure, the solution of Manko *et al.* (2000) has been compared by Stute and Camenzind (2002) and by Berti and Stergioulas (2004) with the numerical solutions for NSs calculated by Cook *et al.* (1994) and with those derived by Berti and Stergioulas (2004), respectively. However, being a generalization of the solution of Tomimatsu and Sato (1972), it cannot describe slowly rotating compact stars (see e.g. Berti and Stergioulas 2004), but the dynamics of astrophysical objects with anisotropic stresses (see Dubeibe *et al.* 2007, for details).

Following a similar procedure, based on tests of the ISCOs radii on the equatorial plane of the rotating neutron stars ob-

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tained by [Berti and Stergioulas \(2004\)](#), it has been shown that the six-parametric solution of [Pachón et al. \(2006\)](#) (hereafter PRS solution, see Sec. 2 for details) is more accurate than the model of [Manko et al. \(2000\)](#). In addition, being a generalization of the Kerr solution, this solution can be used for arbitrary rotation rates.

Besides the ISCOs radii, there are additional physical properties that can be computed with analytic and numerical models and thus useful to compare and contrast the accuracy of analytic exact models. The aim of this article is to analyze the properties of orbital frequencies of neutral test particles in the PRS and in the Kerr geometries with especial focus on the Keplerian ν_K , frame-dragging (Lense-Thirring) ν_{LT} , as well as the precession (oscillation) frequencies of the radial and vertical motions, ν_ρ^P (ν_ρ^{OS}) and ν_z^P (ν_z^{OS}), respectively.

The relevance of these frequencies relies on the fact that they are often invoked to explain the Quasi-Periodic Oscillations (QPOs) observed in some relativistic astrophysical systems such as Low Mass X-ray Binaries (LMXBs), binary systems harboring either a NS or a black hole (BH) accreting matter from a companion star. For instance, within the Relativistic Precession Model (RPM) introduced by [Stella and Vietri \(1998\)](#); [Morsink and Stella \(1999\)](#); [Stella et al. \(1999\)](#); [Stella and Vietri \(1999\)](#), the kHz QPOs are interpreted as a direct manifestation of the modes of relativistic epicyclic motion of blobs arising at various radii r in the inner parts of the accretion disk around the compact object (see Sec. 6, for details).

In addition to the RPM, the Keplerian, precession and oscillation frequencies are used in other QPO theoretical models (see e.g. [Lin et al. 2011](#), for a recent comparison of the existing models). Due to the influence of general relativistic effects in the determination of such frequencies, an observational confirmation of any of the models might lead to an outstanding test of general relativity in the strong field regime. In this line, it is of interest to compare and contrast the orbital frequencies given by the Kerr solution and by the PRS solution (see Sec. 3), which help to establish the differences between possible BH and NS signatures. We emphasize in this article the major role of the quadrupole moment as well as of the octupole moment of the object, whose possible measurement can be used as a tool to test the no-hair theorem of BHs (see e.g. [Johannsen and Psaltis 2011](#)) and to discriminate between the different theoretical models proposed to explain the physics at interior and exterior of the Neutron Stars. Additionally, in the case of NSs, the interpretation of QPOs as the manifestation of orbital motion frequencies might lead to crucial information of the NS parameters such as mass, angular momentum (see e.g. [Stella and Vietri 1998](#); [Török et al. 2010](#)), and quadrupole moment (see e.g. [Morsink and Stella 1999](#)). These parameters reveal, at the same time, invaluable information about the EoS of nuclear matter.

The article is organized as follows. In Sec. 2 we recall the properties of the PRS solution. The computation of the orbital frequencies as well as the comparison of their features in the Kerr and in the PRS spacetimes, is shown in Sec. 3. In Sec. 4 we study the accuracy of the analytic formulas of the periastron and nodal frequencies derived by [Ryan \(1995\)](#) for stationary axially symmetric spacetimes. In Sections 5 and 6 we discuss the accuracy of the PRS solution in describing the frequencies of realistic NS models and its relevance in the Relativistic Precession Model, respectively. The conclusions of this work and a discussion on possible additional effects to

be accounted for in the determination of the orbital frequencies, e.g. the effect of magnetic dipole moment, are outlined in Sec. 7.

2. THE PRS ANALYTIC EXACT SOLUTION

We first recall the PRS analytic model ([Pachón et al. 2006](#)), for the exterior gravitational field of a compact object⁴. In the stationary axisymmetric case, the simplest form of the metric can be written as ([Papapetrou 1953](#))

$$ds^2 = -f(dt - \omega d\phi)^2 + f^{-1} \left[e^{2\gamma}(d\rho^2 + dz^2) + \rho^2 d\phi^2 \right], \quad (1)$$

where f , ω and γ are functions of the quasi-cylindrical Weyl coordinates (ρ, z) . Thus, the components of the metric tensor $g_{\mu\nu}$ are

$$g_{\phi\phi} = \frac{\rho^2}{f(\rho, z)} - f(\rho, z)\omega(\rho, z)^2, \quad (2)$$

$$g_{tt} = -f(\rho, z), \quad (3)$$

$$g_{t\phi} = f(\rho, z)\omega(\rho, z), \quad (4)$$

$$g_{zz} = g_{\rho\rho} = \frac{e^{2\gamma(\rho, z)}}{f(\rho, z)} = \frac{1}{g^{zz}} = \frac{1}{g^{\rho\rho}}. \quad (5)$$

Using the above line element, the Einstein-Maxwell equations can be reformulated, via Ernst's procedure in terms of two complex potentials $\mathcal{E}(\rho, z)$ and $\Phi(\rho, z)$ ([Ernst 1968a,b](#)). By means of Sibgatullin's integral method ([Sibgatullin 1991](#); [Manko and Sibgatullin 1993](#)) this system of equations can be solved via

$$\mathcal{E}(z, \rho) = \int_{-1}^1 \frac{d\sigma}{\pi} \frac{e(\xi)\mu(\sigma)}{\sqrt{1-\sigma^2}}, \quad (6)$$

$$\Phi(z, \rho) = \int_{-1}^1 \frac{d\sigma}{\pi} \frac{f(\xi)\mu(\sigma)}{\sqrt{1-\sigma^2}}, \quad (7)$$

where $e(z) := \mathcal{E}(z, \rho = 0)$ and $f(z) := \Phi(z, \rho = 0)$. The unknown function $\mu(\sigma)$ must satisfy the singular integral equation

$$\oint_{-1}^1 \frac{\mu(\sigma)[e(\xi) + \tilde{e}(\eta) + 2f(\xi)\tilde{f}(\eta)]d\sigma}{(\sigma - \tau)\sqrt{1-\sigma^2}} = 0 \quad (8)$$

and the normalizing condition

$$\int_{-1}^1 \frac{\mu(\sigma)d\sigma}{\sqrt{1-\sigma^2}} = \pi, \quad (9)$$

where $\xi = z + i\rho\sigma$, $\eta = z + i\rho\tau$, ρ and z being the Weyl-Papapetrou quasi-cylindrical coordinates, $\sigma, \tau \in [-1, 1]$, $\tilde{e}(\eta) := \overline{e(\bar{\eta})}$, $\tilde{f}(\eta) := \overline{f(\bar{\eta})}$ and the overbar stands for complex conjugation. In ([Pachón et al. 2006](#)), the Ernst potentials were chosen as

$$e(z) = \frac{z^3 - z^2(m + ia) - kz + is}{z^3 + z^2(m - ia) - kz + is}, \quad (10)$$

$$f(z) = \frac{qz^2 + i\mu z}{z^3 + z^2(m - ia) - kz + is}.$$

⁴ Mathematica 8.0 scripts with the solution, some limiting cases as well as the the calculations presented in this paper are available at <http://www.chem.utoronto.ca/~lpachon/scripts/nstars>

We calculate the multipole moments following the procedure of [Hoenselaers and Perjes \(1990\)](#). We denote the mass multipoles by M_i while, the current (rotation) multipoles, by S_i . The electric multipoles are denoted by Q_i and the magnetic ones by \mathcal{B}_i . Thus, for the PRS solution we have

$$\begin{aligned} M_0 &= m, & M_2 &= mk - ma^2, & \dots \\ S_1 &= ma, & S_3 &= -ma^3 + 2mak - ms, & \dots \\ Q_0 &= q, & Q_2 &= -a^2q - a\mu + kq, & \dots \\ \mathcal{B}_1 &= \mu + aq, & \mathcal{B}_3 &= -a^2\mu + \mu k - a^3q + 2akq - qs, & \dots \end{aligned} \quad (11)$$

This allows us to identify m as the total mass, a as the total angular momentum per unit mass ($a = J/m$, being J the total angular momentum); while k , s , q and μ are associated to the mass-quadrupole moment M_2 , current octupole S_3 , electric charge and magnetic dipole, respectively.

The potentials (10) can be written in an alternative way, we mean

$$e(z) = 1 + \sum_{i=3}^3 \frac{e_i}{z - \beta_i}, \quad f(z) = \sum_{i=3}^3 \frac{f_i}{z - \beta_i}, \quad (13)$$

with

$$e_j = (-1)^j \frac{2m\beta_j^2}{(\beta_j - \beta_k)(\beta_j - \beta_i)}, \quad (14)$$

$$f_j = (-1)^{j+1} \frac{i\mu + d\beta_j}{(\beta_j - \beta_k)(\beta_j - \beta_i)}, \quad i, k \neq j. \quad (15)$$

Then, using Eqs. (6) and (10), we obtain the Ernst potentials

$$\mathcal{E} = \frac{A+B}{A-B}, \quad \Phi = \frac{C}{A-B}, \quad (16)$$

and the metric functions in the whole spacetime

$$\begin{aligned} f &= \frac{A\bar{A} - B\bar{B} + C\bar{C}}{(A-B)(\bar{A} - \bar{B})}, & e^{2\gamma} &= \frac{A\bar{A} - B\bar{B} + C\bar{C}}{K\bar{K} \prod_{n=1}^6 r_n}, & (17) \\ \omega &= \frac{\text{Im}[(A+B)\bar{H} - (\bar{A} + \bar{B})G - C\bar{I}]}{A\bar{A} - B\bar{B} + C\bar{C}}, & & & (18) \end{aligned}$$

where the functions A , B , C , H , G , K , and I can be found in the Appendix A.

The PRS electrovacuum exact solution belongs to the extended N -soliton solution of the Einstein-Maxwell equations derived by [Ruiz et al. \(1995\)](#), in the particular case $N = 3$. In addition, the functional form of the metric functions resembles the one derived previously by [Bretón et al. \(1999\)](#). Besides the limiting cases discussed in [Pachón et al. \(2006\)](#) it is worth mentioning that, in the vacuum case $q = 0$ and $\mu = 0$, for $s = 0$ this solution reduces to the solution of [Manko et al. \(1995\)](#) under the same physical conditions, namely $q = 0$, $c = 0$ and $b = 0$ in [Manko et al. \(1995\)](#).

3. ORBITAL MOTION FREQUENCIES ON THE EQUATORIAL PLANE

Although for the case of compact stars contributions from the magnetic field could be relevant (see e.g. [Bakala et al. 2010](#); [Sanabria-Gómez et al. 2010](#); [Bakala et al. 2012](#)), we focus in this work on the frequencies of neutral particles orbiting a neutral compact object. We calculate here the Keplerian $\nu_K = \Omega_K/(2\pi)$, frame-dragging (Lense-Thirring) $\nu_{LT} =$

$\Omega_{LT}/(2\pi)$, radial oscillation and precession, $\nu_\rho^{\text{OS}} = \Omega_\rho^{\text{OS}}/(2\pi)$ and $\nu_\rho^{\text{P}} = \Omega_\rho^{\text{P}}/(2\pi)$, and vertical oscillation and precession frequencies, $\nu_z^{\text{OS}} = \Omega_z^{\text{OS}}/(2\pi)$ and $\nu_\rho^{\text{P}} = \Omega_\rho^{\text{P}}/(2\pi)$, respectively.

The geodesic motion of test particles along the radial coordinate, on the equatorial plane $z = 0$, is governed by the effective potential (see e.g. [Ryan 1995](#))

$$V(\rho) = 1 - \frac{E^2 g_{\phi\phi} + 2ELg_{t\phi} + L^2 g_{tt}}{g_{t\phi}^2 - g_{tt}g_{\phi\phi}}, \quad (19)$$

where, for circular orbits, the energy E and angular momentum L are determined by the conditions $V = 0$ and $dV/d\rho = 0$ (see Eqs. 22–23). The frequencies at the ISCO's location (determined by the additional condition $d^2V/d\rho^2 = 0$) are of particular interest. Thus, before starting the discussion of the frequencies, it is important to explore the ISCO parametric dependence. We report here, as standard in the literature, the *physical* ISCO radius given by $\sqrt{g_{\phi\phi}}$ evaluated at the root of Eq. (19) that gives the coordinate ISCO radius. In the upper panel of Fig. 1 we plotted contours of constant ISCO radii as a function of the dimensionless angular momentum parameter $j = J/M_0^2$ and the star quadrupole moment M_2 , for the PRS solution. The use of the dimensionless parameter j in the horizontal axis allows to, qualitatively, relate deviations of the contour lines from vertical lines to the influence of the quadrupole moment. We can see that the ISCO radius decreases for increasing j and decreasing M_2 . A quantitative measurement of this influence could be derived from the effective slope of the contour lines. We are interested in the comparison with the Kerr geometry, so in the lower panel, we plotted contours of constant ratio $r_{\text{ISCO,PRS}}/r_{\text{ISCO,Kerr}}$ as a function of j and the difference between the quadrupole moment of the PRS solution $M_{2,\text{PRS}}$ and the Kerr quadrupole $M_{2,\text{Kerr}} = -ma^2$, i.e. $M_{2,\text{PRS}} - M_{2,\text{Kerr}} = M_{2,\text{PRS}} + ma^2 = mk$, see Eq. (11). Deviations from the Kerr geometry are evident. Negative values of the angular momentum correspond to the radii of the counter-rotating orbits obtained here through the change $g_{t\phi} \rightarrow -g_{t\phi}$ (see discussion below).

We stress that the accuracy of the PRS solution for describing the ISCO radius of *realistic* NSs was already shown to be higher with respect to other analytic models (see [Pachón et al. 2006](#), for details). In Table 1 we compare the ISCO radius for two rapidly rotating NS, models 20 and 26, of Table VI of [Pappas and Apostolatos \(2012\)](#) for the EoS L. The lowest multipole moments of the analytic models are fixed to the numerical values obtained by [Pappas and Apostolatos \(2012\)](#). In the case of the Kerr solution, only M_0 and J can be fixed, while M_2 , and S_3 have values that depend on M_0 and J and therefore cannot be fixed. For the PRS solution with $s = 0$, M_0 , J and M_2 can be fixed while S_3 remains induced by the lower moments. We present also the ISCO radius obtained by fixing M_0 , J , M_2 , as well as S_3 in the PRS analytic exact model.

In Figs. 1–6, we have fixed as an example $M_0 = m = 1.88M_\odot = 2.78$ km, and $s = 0$. We recall that the quadrupole moment in the geometric units used here (km^3) is related to the one in CGS units by $M_2^{\text{CGS}} = (10^{15}c^2/G)M_2^{\text{geo}} = 1.35 \times 10^{43}(M_2^{\text{geo}}/\text{km}^3)$ g cm², and the mass of the Sun is $M_\odot^{\text{geo}} = 1.477$ km. The dimensionless angular momentum j is obtained from the CGS values of J and M_0 as $j = cJ/(GM_0^2)$.

It is appropriate to compare the range of values of $j = J/M_0^2$ and M_2 used in Figs. 1–6 with typical values of a NS. For the used mass $M_0 = 1.88M_\odot$, [Morsink and Stella \(1999\)](#) ob-

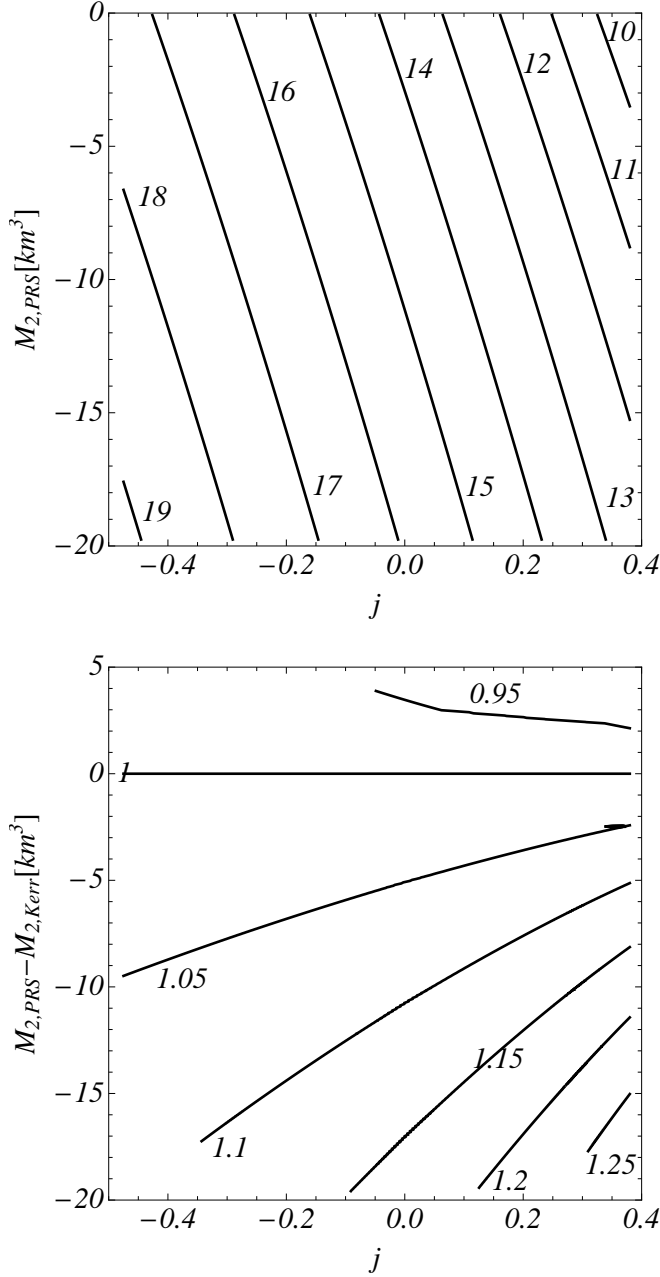


FIG. 1.— Upper panel: Contours of constant ISCO radius as a function of the dimensionless angular momentum parameter $j = J/M_0^2$ and the quadrupole moment M_2 for the PRS solution, for a compact object with mass $M_0 = m = 1.88M_\odot = 2.78$ km. Contours are labeled by the corresponding value of the ISCO radius in km. Negative values of j depict the counter-rotating case and negative values of the quadrupole moment M_2 correspond to oblate configurations. The values of M_2 are in the range $0 \leq M_2 \leq 20$ km³ that corresponds in CGS units to $0 \leq M_2 \leq -2.7 \times 10^{44}$ g cm², which covers the typical range of fast rotating NSs. Lower panel: Contours of constant ratio $r_{\text{ISCO,PRS}}/r_{\text{ISCO,Kerr}}$ as a function of j and the difference $M_{2,\text{PRS}} - M_{2,\text{Kerr}}$. The quadrupole moment difference is comprised in the range $-2.7 \times 10^{44} \leq M_2 \leq 6.8 \times 10^{43}$ g cm².

tained a quadrupole moment $M_2 = -5.3 \times 10^{43}$ g cm² = 3.93 km³, with the latter value in geometric units, for a NS of angular rotation frequency $\nu_s = 290$ Hz (rotation period of 3.45 milliseconds), corresponding to a dimensionless angular momentum $j = J/M_0^2 = 0.19$, for the EoS L. For a fixed mass the quadrupole moment is an increasing func-

	R_N [km]	R_{SS} [km]	R_{Kerr} [km]	$R_{\text{PRS},s=0}$ [km]	R_{PRS} [km]
M20	19.81	13.39	16.14	19.28	18.99
M26	19.87	17.16	15.94	19.65	19.54

TABLE 1

COMPARISON OF THE ISCO RADIUS FOR THE SELECTED NS MODELS 20 AND 26 OF TABLE VI OF PAPPAS AND APOSTOLATOS (2012) FOR THE EoS L. MODEL 20: $M_0 = 4.167$ km ($2.82M_\odot$), $j = J/M_0^2 = 0.70$, $M_2 = -79.8$ km³ (-1.08×10^{45} g cm²) AND $S_3 = -401.0$ km⁴. MODEL 26: $M_0 = 4.36$ km ($2.95M_\odot$), $j = J/M_0^2 = 0.56$, $M_2 = -45.2$ km³ (-6.10×10^{44} g cm²) AND $S_3 = -170.0$ km⁴. THE SUBSCRIPT N STANDS FOR THE NUMERICAL CALCULATION OF PAPPAS AND APOSTOLATOS (2012) AND SS STANDS FOR THE SHIBATA AND SASAKI (1998) APPROXIMATED ISCO RADIUS EXPRESSION.

tion of j because an increasing of the angular momentum at fixed mass results in an increasing of the oblateness (eccentricity) of the star, and so the quadrupole moment. Based on this fact, it is clear that not all the values of the (M_2, j) pairs of quadrupole and angular momentum depicted in, e.g., Fig. 1 are physically meaningful. The maximum rotation rate of a neutron star taking into account both the effects of general relativity and deformations has been found to be $\nu_{s,\text{max}} = 1045(M_0/M_\odot)^{1/2}(10\text{ km}/R)^{3/2}$ Hz, largely independent on the EoS (see Lattimer and Prakash 2004, for details). Corresponding to this maximum rotation rate, the angular momentum is $J_{\text{max}} = 2\pi\nu_{s,\text{max}}I \sim 6.56 \times 10^{48}I_{45}$ g cm² s⁻¹, and $j_{\text{max}} = GJ_{\text{max}}/(cM_0^2) \sim 0.74I_{45}/(M_0/M_\odot)^2$, where I_{45} is the moment of inertia of the NS in units of 10^{45} g cm². The fastest observed pulsar is PSR J1748-2246ad with a rotation frequency of 716 Hz (Hessels et al. 2006), which constrains the mass of the NS to $M_0 \geq 0.47(R/10\text{ km})^3M_\odot$, and $j \sim 0.51I_{45}/(M_0/M_\odot)^2$, which becomes $j \sim 0.26I_{45}$ for a canonical NS of $M_0 = 1.4M_\odot$.

3.1. Keplerian Frequency

Now we turn into the frequencies analysis. For stationary axially symmetric spacetimes, the frequency of Keplerian orbits is given by (see e.g. Ryan 1995)

$$\Omega_K = \frac{-g_{t\phi,\rho} \pm \sqrt{g_{t\phi,\rho}^2 - g_{\phi\phi,\rho}g_{t\rho,\rho}}}{g_{\phi\phi,\rho}}, \quad (20)$$

where a colon stands for partial derivative with respect to the indicated coordinate and ‘+’ and ‘-’ stands for corotating and counter-rotating orbits, respectively.

For the case of static spacetimes, i.e. for $\omega = 0$ and therefore $g_{t\phi} = 0$, $\Omega_K = \pm \sqrt{-g_{\phi\phi,\rho}g_{t\rho,\rho}}/g_{\phi\phi,\rho}$ and the energy E and angular momentum L per mass μ of the test particle can be expressed in terms of the metric tensor components (see e.g. Ryan 1995),

$$\frac{E}{\mu} = \frac{-g_{tt}}{\sqrt{-g_{tt} - g_{\phi\phi}\Omega_K^2}}, \quad \frac{L}{\mu} = \frac{g_{\phi\phi}\Omega_K}{\sqrt{-g_{tt} - g_{\phi\phi}\Omega_K^2}}. \quad (21)$$

From here, it is clear that taking the negative branch of the root for Ω_K in Eq. (20) is equivalent to studying a particle with opposite angular momentum, i.e. $L_{\text{count-rot}} = -L_{\text{co-rot}}$. Thus, in the static case the magnitude of the energy and angular momentum are invariant under the change $\Omega_K \rightarrow -\Omega_K$.

Now we consider the case of stationary space times, $\omega \neq 0$. The energy E and angular momentum L per mass μ are, in this

case, given by (see e.g. [Ryan 1995](#))

$$\frac{E}{\mu} = \frac{-g_{tt} - g_{t\phi}\Omega_K}{\sqrt{-g_{tt} - 2g_{t\phi}\Omega_K - g_{\phi\phi}\Omega_K^2}}, \quad (22)$$

$$\frac{L}{\mu} = \frac{g_{t\phi} + g_{\phi\phi}\Omega_K}{\sqrt{-g_{tt} - 2g_{t\phi}\Omega_K - g_{\phi\phi}\Omega_K^2}}. \quad (23)$$

The counter-rotating condition expressed by the negative branch of Eq. (20), can be generated by the change $g_{t\phi} \rightarrow -g_{t\phi}$, which seems to be a more physical and transparent condition. In contrast to the static case, the counter-rotating orbit has now different energy and different magnitude of the angular momentum due the presence of the dragging of inertial frames, characterized by the metric component $g_{t\phi}$ (cf. Eq. (26) below). In a nutshell, the dynamics of counter-rotating orbits of a test-particle can be derived, starting from the positive branch of Eq. (20), by considering a spacetime with $g_{t\phi} \rightarrow -g_{t\phi}$.

For the vacuum case, a similar analysis as the one developed by [Herrera et al. \(2006\)](#), clearly shows that the change in the global sign of $g_{t\phi}$ is achieved by changing not only the angular momentum of the star, $J \rightarrow -J$, but *all the rotational multipolar moments*. For the Kerr metric this change is obtained by changing the sign of the parameter a (see Appendix B) while in the PRS solution we need additionally change the sign of the parameter s associated to *differential rotation*, i.e., by changing $a \rightarrow -a$ and $s \rightarrow -s$ ⁵.

Once we have clarified this important issue about the co-rotating and counter-rotating orbits, we proceed to analyze the functional dependence of the Keplerian frequency on the multipole moments. In the upper panel of Fig. 2 we plotted contours of constant Keplerian frequency for the PRS solution, $\nu_{K,PRS} = \Omega_{K,PRS}/(2\pi)$, as a function of the dimensionless angular momentum parameter j and the quadrupole moment $M_{2,PRS}$, at the ISCO radius. It can be seen that the influence of the quadrupole moment is non-negligible, as evidenced from the departure of the contour lines from vertical lines. The Keplerian frequency grows for increasing J and M_2 . In the lower panel, we plotted contours of constant ratio $\nu_{K,PRS}/\nu_{K,Kerr}$ as a function of j and the difference between the quadrupole moment of the PRS solution, $M_{2,PRS}$, and the Kerr quadrupole, $M_{2,Kerr}$.

It is appropriate to recall here that because the Keplerian as well as the other frequencies calculated below are evaluated using formulas in the coordinate frame, see for instance Eq. (20), they must be evaluated at coordinate radii ρ and not at physical radii given by $\sqrt{g_{\phi\phi}}$. In the specific case of the ISCO the frequencies are evaluated at the radius that simultaneously solves the equations $V = 0$, $dV/d\rho = 0$, and $d^2V/d\rho^2 = 0$, where V is the effective potential (19).

3.2. Oscillation and Precession Frequencies

The radial and vertical oscillation (or epicyclic) frequencies are the frequencies at which the periastron and orbital plane of a circular orbit oscillates if we apply slightly radial and vertical perturbations to it, respectively. According to [Ryan \(1995\)](#), in stationary axially symmetric vacuum spacetimes described by the Weyl-Papapetrou metric (1), the radial and

⁵ For the vacuum case, in the solution by [Manko et al. \(2000\)](#), the sign change of $g_{t\phi}$ is obtained after performing simultaneously the replacements $a \rightarrow -a$ and $b \rightarrow -b$.

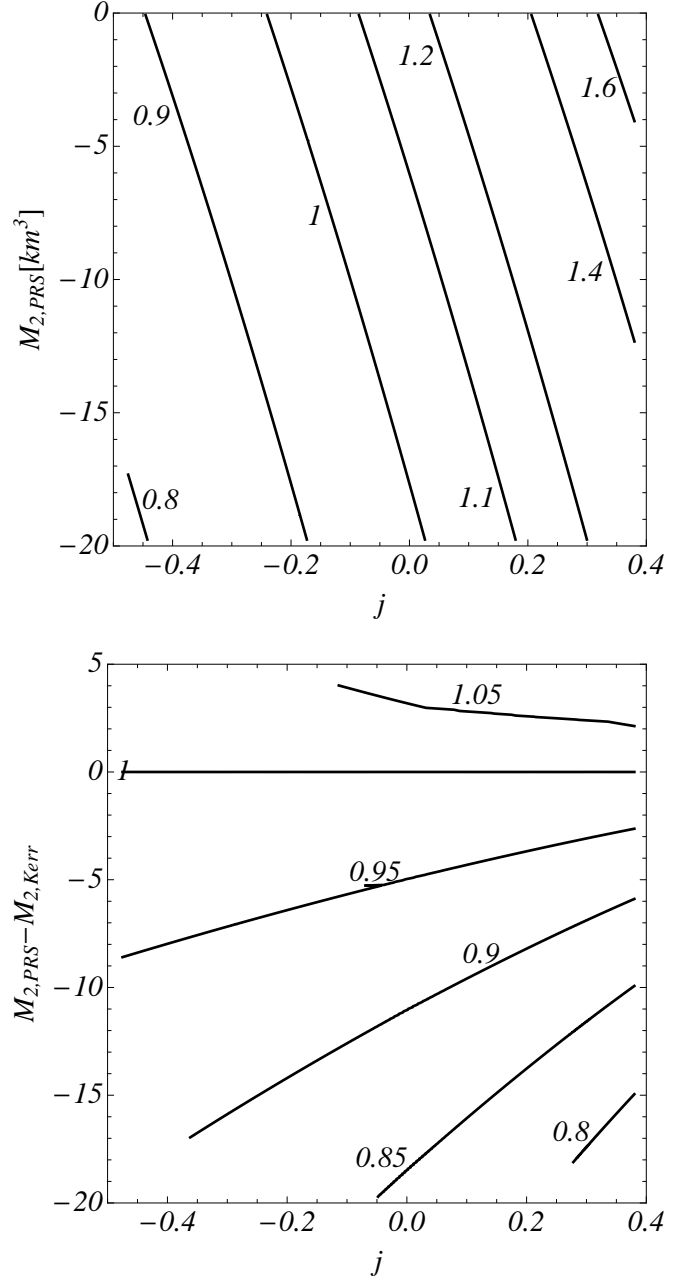


Fig. 2.— Upper panel: Contours of constant ν_K (in kHz) as a function of the the dimensionless angular momentum parameter $j = J/M_0^2$ and quadrupole moment M_2 for the PRS solution, at the ISCO radius, for a compact object with mass $M_0 = m = 1.88M_\odot = 2.78$ km. Lower panel: Contours of constant ratio $\nu_{K,PRS}/\nu_{K,Kerr}$ as a function of j and the difference $M_{2,PRS} - M_{2,Kerr}$, at the ISCO radius.

vertical epicyclic frequencies can be obtained as

$$\begin{aligned} \nu_\alpha^{\text{OS}} = \frac{1}{2\pi} \left\{ -\frac{g^{\alpha\alpha}}{2} \left[(g_{tt} + g_{t\phi}\Omega)^2 \left(\frac{g_{\phi\phi}}{\rho^2} \right)_{,\alpha\alpha} \right. \right. \\ - 2(g_{tt} + g_{t\phi}\Omega)(g_{t\phi} + g_{\phi\phi}\Omega) \left(\frac{g_{t\phi}}{\rho^2} \right)_{,\alpha\alpha} \\ \left. \left. + (g_{t\phi} + g_{\phi\phi}\Omega)^2 \left(\frac{g_{tt}}{\rho^2} \right)_{,\alpha\alpha} \right] \right\}^{1/2}, \quad (24) \end{aligned}$$

and the corresponding periastron (ν_ρ^P) and nodal (ν_z^P) precession frequencies as

$$\nu_\alpha^P = \nu_K - \nu_\alpha^{\text{OS}}. \quad (25)$$

where $\alpha = \{\rho, z\}$, respectively, and $\nu_K = \Omega_K/(2\pi)$ is the Keplerian orbital frequency with Ω_K given by Eq. (20).

In the upper panel of Fig. 3, we plotted contours of constant nodal precession frequency ν_z^P at the ISCO radius as a function of $j = J/M_0^2$ and M_2 for the PRS solution, at the ISCO radius. We can see now that the influence of the quadrupole moment is quite important. The nodal precession frequency increases for increasing J and decreasing M_2 , at fixed M_0 . In the lower panel we plotted contours of constant ratio $\nu_{z,\text{PRS}}^P/\nu_{z,\text{Kerr}}^P$, at the ISCO radius, as a function of j and the difference $M_{2,\text{PRS}} - M_{2,\text{Kerr}}$, in order to evidenciate deviations from the Kerr solution. The radial oscillation frequency ν_ρ^{OS} vanishes at the ISCO radius and therefore at such location the radial precession frequency equals the Keplerian frequency, whose contours have been plotted in Fig. 2.

In Figs. 4 and 5 we plotted the nodal precession frequency ν_z^P and the radial oscillation frequency ν_ρ^{OS} as a function of the Keplerian frequency ν_K , respectively, for both the Kerr and PRS solutions. As an example, we have shown the results for rotating NS models 20 and 26 of Table VI of Pappas and Apostolatos (2012), for the EoS L. The lowest multipole moments of the PRS solution M_0 , J , M_2 , and S_3 have been fixed to the numerical values obtained by Pappas and Apostolatos (2012). In the case of the Kerr solution, only M_0 and J can be fixed, while M_2 , and S_3 have values induced by the lower moments M_0 and J . For the PRS solution with $s = 0$, M_0 , J and M_2 can be fixed while S_3 cannot be fixed and depends on the lower moments. The results for the PRS analytic model obtained by fixing M_0 , J , M_2 , as well as S_3 are also shown.

The deviations of the quadrupole and current octupole moments given by the Kerr solution from the numerical values of Pappas and Apostolatos (2012) can be used to show the low accuracy of the Kerr solution to describe fast rotating NSs. The accuracy of the PRS solution in describing the ISCO radii of these two models has been shown in Table 1 of Section 3.

In Figs. 4 and 5 we can see the differences of the $\nu_z^P - \nu_K$ and $\nu_\rho^{\text{OS}} - \nu_K$ relations between the Kerr and PRS solutions for realistic NS models. The deviations of the Kerr solution, especially at fast rotation rates, are evident because of the influence of the deformation (quadrupole M_2) of the star as well as, although in less proportion, of the octupole current S_3 . In general, we observe that the larger the angular momentum, the poorer the performance of the predictions of Kerr solution.

We have also shown in Figs. 4–5 the influence of the current octupole S_3 in the determination of the precession and oscillation frequencies. We found that the effect of S_3 is only appreciable for the fastest models. The minor influence, in this case, of the current octupole S_3 is expected from the small values of the parameter s needed to fit the numerical values of Pappas and Apostolatos (2012). Clearly, larger values of the parameter s needed to fit realistic values of S_3 will enhance as well deviations from the Kerr spacetime.

The effects of a multipolar structure that deviates from the one of the Kerr geometry on the various quantities analyzed here are relevant for instance in the RPM of the QPOs observed in LMXBs (see e.g. Stella and Vietri (1998); Morsink and Stella (1999); Stella et al. (1999);

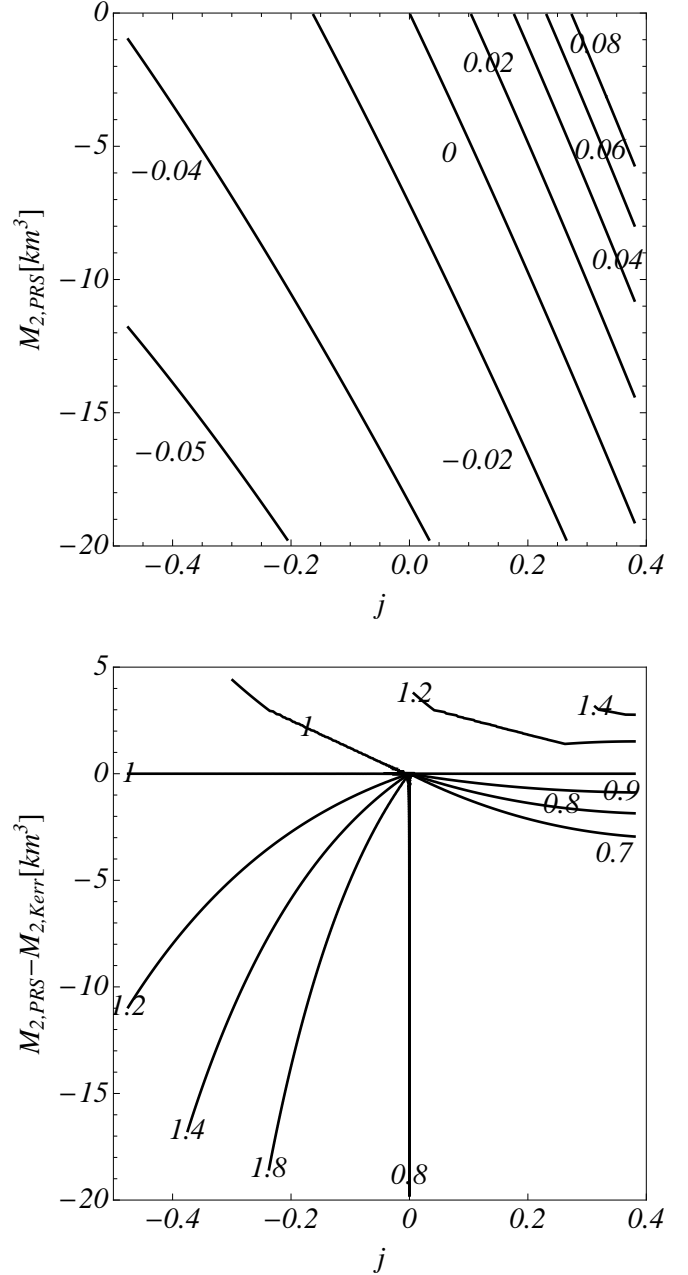


FIG. 3.— Upper panel: ν_z^P (in kHz) as a function of the dimensionless angular momentum parameter $j = J/M_0^2$ and quadrupole moment M_2 for the PRS solution, at the ISCO radius, for a compact object with mass $M_0 = m = 1.88M_\odot = 2.78$ km. Lower panel: Contours of constant ratio $\nu_{z,\text{PRS}}^P/\nu_{z,\text{Kerr}}^P$ as a function of j and the difference $M_{2,\text{PRS}} - M_{2,\text{Kerr}}$, at the ISCO radius.

Stella and Vietri (1999) and Section 6, for details).

3.3. Dragging of Inertial Frames

It is known that a prediction of general relativity is that a rotating object makes a zero angular momentum test particle to orbit around it, namely it *drags* the particle into the direction of its rotation angular velocity; such an effect is called dragging of inertial frames or Lense-Thirring effect. Consequently, oblique particle orbit planes with respect to the source equatorial plane will precess around the rotation axis of the object. In stationary axially symmetric spacetimes described by the

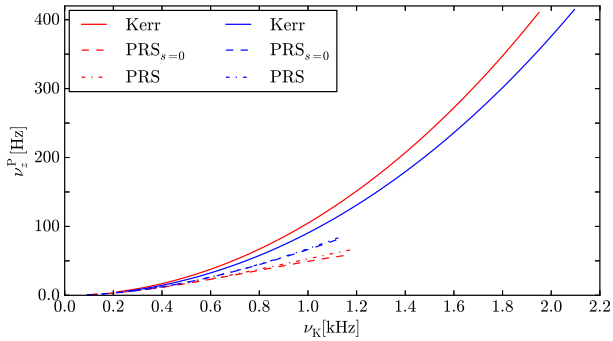


FIG. 4.— Nodal precession frequency ν_z^P versus Keplerian frequency ν_K given by the Kerr and PRS analytic solutions. The lowest multipole moments have been fixed from the rotating NS models 20 (red curves) and 26 (blue curves) of the Table VI of Pappas and Apostolatos (2012) for the EoS L. Model 20: $M_0 = 4.167$ km ($2.82M_\odot$), $j = J/M_0^2 = 0.70$, $M_2 = -79.8$ km³ (-1.08×10^{45} g cm²) and $S_3 = -401.0$ km⁴. Model 26: $M_0 = 4.36$ km ($2.95M_\odot$), $j = J/M_0^2 = 0.56$, $M_2 = -45.2$ km³ (-6.10×10^{44} g cm²) and $S_3 = -170.0$ km⁴.

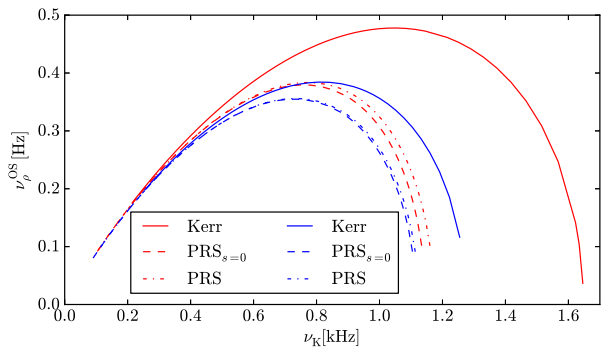


FIG. 5.— Radial oscillation frequency ν_ρ^{OS} versus Keplerian frequency ν_K given by the Kerr and PRS analytic solutions. The lowest multipole moments has been fixed from the rotating NS models 20 (red curves) and 26 (blue curves) of the Table VI of Pappas and Apostolatos (2012) for the EoS L. Model 20: $M_0 = 4.167$ km ($2.82M_\odot$), $j = J/M_0^2 = 0.70$, $M_2 = -79.8$ km³ (-1.08×10^{45} g cm²) and $S_3 = -401.0$ km⁴. Model 26: $M_0 = 4.36$ km ($2.95M_\odot$), $j = J/M_0^2 = 0.56$, $M_2 = -45.2$ km³ (-6.10×10^{44} g cm²) and $S_3 = -170.0$ km⁴.

metric (1) the frame dragging precession frequency is given by (see e.g. Ryan 1995)

$$\nu_{LT} = -\frac{1}{2\pi} \frac{g_{t\phi}}{g_{\phi\phi}}. \quad (26)$$

Many efforts have been done to test the predictions of general relativity around the Earth such as the analysis of the periastron precession of the orbits of the Laser Geodynamics Satellites, LAGEOS and LAGEOS II, (see e.g. Lucchesi and Peron 2010) and the relativistic precession of the gyroscopes on-board the Gravity Probe B satellite (see Everitt et al. 2011, for details). The latter experiment measured a frame dragging effect within an accuracy of 19% with respect to the prediction of general relativity.

The smallness of this effect around the Earth makes such

measurements quite difficult and has represented a multi year challenge for Astronomy. The frame dragging precession increases with the increasing of the angular momentum of the rotating object and therefore a major hypothetical arena for the searching of more appreciable Lense-Thirring precession is the spacetime around compact objects such as BHs and NSs. The much stronger gravitational field of these objects with respect to the Earth one allow them to attain much faster angular rotation rates and so larger angular momentum.

Stella and Vietri (1998) showed how, in the weak field slow rotation regime, the vertical precession frequency ν_z^P (orbital plane precession frequency) can be divided into one contribution due to the Lense-Thirring precession and another one due to the deformation (non-zero quadrupole moment) of the rotating object, both of them comparable from the quantitative point of view. These frequencies could be in principle related to the motion of the matter in the accretion disks around BHs and NSs and thus particularly applicable to LMXBs. For fast rotating NSs and BHs the frequency at which the orbital plane, and so the frame dragging precession frequency, can reach values of the order of tens of Hz (see e.g. Stella and Vietri (1998) and Figs. 3 and 4).

Thus, it is clear that an observational confirmation of the relativistic precession of matter around either a NS or a BH will lead to an outstanding test of the general relativity in the strong field regime and, at the same time, an indirect check of the large effects of the frame dragging in the exterior spacetime of compact objects (see e.g. Morsink and Stella 1999, for details).

Although making independent measurements of the frame dragging effect around BHs and NSs is a very complicate task, it is important to know the numerical values of the precession frequency due to the frame dragging with respect to other relativistic precession effects, e.g. geodetic precession. In addition, it is important to know the sensitivity of the precession frequency to the object parameters such as mass, angular momentum, quadrupole, and octupole moment.

In the upper panel of Fig. 6 we plotted contours of constant frame dragging frequency ν_{LT} for the PRS solution, at the ISCO radius, as a function of the the angular momentum per unit mass J/M_0 and the quadrupole moment M_2 , for a compact object mass $M_0 = m = 1.88M_\odot$. Correspondingly, in the lower panel of Fig. 6, we show the differences between the frame dragging precession frequency as predicted by the Kerr and PRS solutions, at the ISCO radius, as a function of $j = J/M_0^2$ and the difference between the quadrupole moments, $M_{2,PRS} - M_{2,Kerr}$.

The influence of the quadrupole moment in the determination of the frame dragging frequency is evident; the frequency ν_{LT} given by a NS is generally smaller than the one given by a BH as can be seen from the value of the ratio $\nu_{LT,PRS}/\nu_{LT,Kerr} < 1$ obtained for configurations with a quadrupole moment that deviates with respect to the one given by the Kerr solution, namely for $M_{2,PRS} - M_{2,Kerr} = M_{2,PRS} + ma^2 = mk \neq 0$, see Eq. (11).

It is also worth mentioning that frame dragging precession can be affected as well by the presence of electromagnetic fields (Herrera et al. 2006, see) and further research in this respect deserves the due attention.

4. ACCURACY OF RYAN'S ANALYTIC FORMULAS

Following a series expansion procedure in powers of $1/\rho$, Ryan (1995) found that the periastron (radial) and nodal (vertical) precession frequencies, ν_ρ^P and ν_z^P given by Eq. (24), can

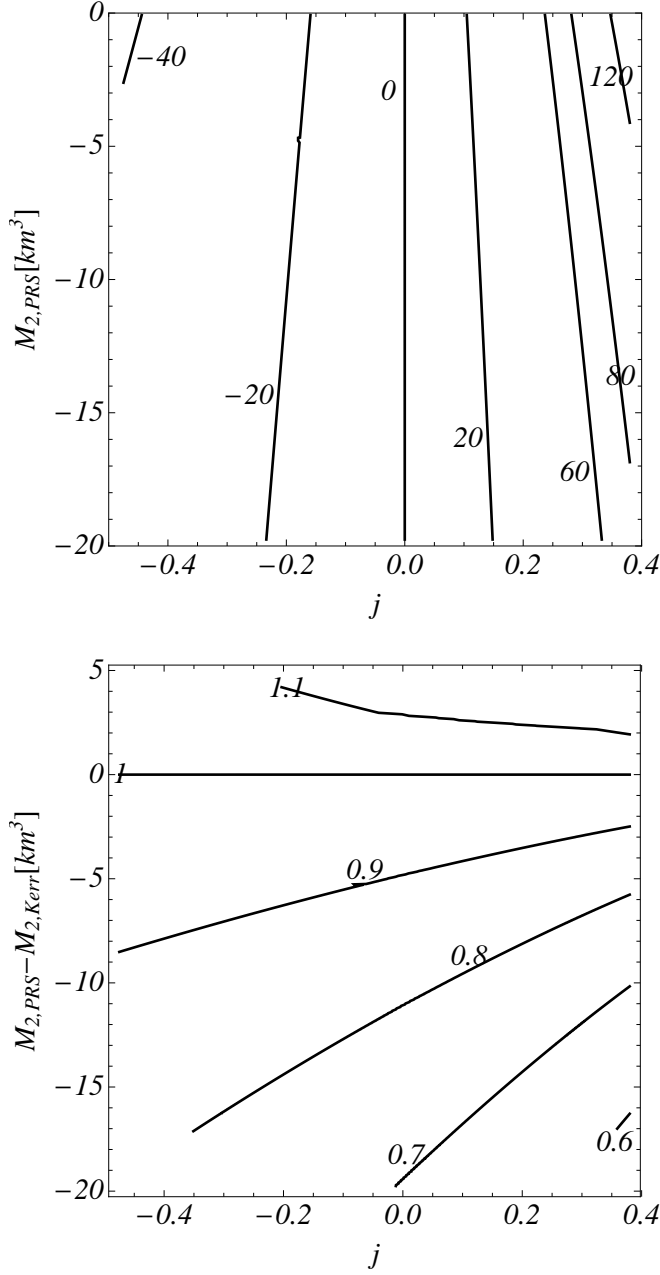


FIG. 6.— Upper panel: Contours of constant ν_{LT} (in Hz) as a function of the angular momentum per unit mass $j = J/M_0^2$ and the quadrupole moment M_2 for the PRS solution, at the ISCO radius, for a compact object with mass $M_0 = m = 1.88M_\odot = 2.78$ km. Lower panel: Contours of constant ratio $\nu_{LT,PRS}/\nu_{LT,Kerr}$ as a function of j and the difference $M_{2,PRS} - M_{2,Kerr}$, at the ISCO radius.

be written as a function of the Keplerian frequency ν_K as

$$\begin{aligned} \frac{\nu_\rho^P}{\nu_K} = & 3\mathcal{V}^2 - 4\frac{S_1}{M_0^2}\mathcal{V}^3 + \left(\frac{9}{2} - \frac{3M_2}{2M_0^3}\right)\mathcal{V}^4 - 10\frac{S_1}{M_0^2}\mathcal{V}^5 \\ & + \left(\frac{27}{2} - 2\frac{S_1^2}{M_0^4} - \frac{21M_2}{2M_0^3}\right)\mathcal{V}^6 + \left(-48\frac{S_1}{M_0^2} - 5\frac{S_1M_2}{M_0^5}\right. \\ & + 9\frac{S_3}{M_0^4}\mathcal{V}^7 + \left(\frac{405}{8} + \frac{2243}{84}\frac{S_1^2}{M_0^4} - \frac{661M_2}{14M_0^3}\right. \\ & \left. - \frac{21M_2^2}{8M_0^6} + \frac{15M_4}{4M_0^5}\right)\mathcal{V}^8 + \dots, \end{aligned} \quad (27)$$

and

$$\begin{aligned} \frac{\nu_z^P}{\nu_K} = & 2\frac{S_1}{M_0^2}\mathcal{V}^3 + \frac{3M_2}{2M_0^3}\mathcal{V}^4 + \left(7\frac{S_1^2}{M_0^4} + 3\frac{M_2}{M_0^3}\right)\mathcal{V}^6 \\ & + \left(11\frac{S_1M_2}{M_0^5} - 6\frac{S_3}{M_0^4}\right)\mathcal{V}^7 + \left(\frac{153}{28}\frac{S_1^2}{M_0^4} + \frac{153M_2}{28M_0^3}\right. \\ & \left. + \frac{39M_2^2}{8M_0^6} - \frac{15M_4}{4M_0^5}\right)\mathcal{V}^8 + \dots, \end{aligned} \quad (28)$$

where $\mathcal{V} = (2\pi M_0 \nu_K)^{1/3}$, $[M_0, M_2, M_4]$ are the lowest three mass moments and, $[S_1, S_3]$, are the lowest two current moments. For the PRS solution in the vacuum case, $M_4 = m(a^4 - 3a^2 + k^2 + 2as)$.

The above formulas are approximate expressions of the periastron and nodal precession frequencies in the weak field (large distances from the source) and slow rotation regimes. We should therefore expect that they become less accurate at distances close to the central object, e.g. at the ISCO radius, and for fast rotating objects. However, such formulas are an important tool to understand the role of the lowest multipole moments on the values of the relativistic precession frequencies, such as the importance of the higher multipole moments at short distances and high frequencies as can be seen from Eqs. (27–28).

At high frequencies, for instance of the order of kHz, deviations from the above scaling laws are appreciable. In Figs. 7 and 8 we compare the radial precession and vertical oscillation frequencies, ν_ρ^P and ν_z^{OS} , as a function of the Keplerian frequency ν_K , as given by the full expressions (24) for the PRS solution and by the approximate formulas (27) and (28), respectively.⁶ The lowest multipole moments M_0 , J , M_2 , and S_3 of the PRS solution have been fixed to the values of two models of the Table VI of Pappas and Apostolatos (2012); Model 2 with $M_0 = 2.071$ km ($1.402M_\odot$), $j = 0.194$, $M_2 = -2.76$ km³ (3.73×10^{43} g cm²), $S_3 = -2.28$ km⁴ and Model 20 with $M_0 = 4.167$ km ($2.82M_\odot$), $j = J/M_0^2 = 0.70$, $M_2 = -79.8$ km³ (-1.08×10^{45} g cm²) and $S_3 = -401.0$ km⁴.

In the $\nu_z^{OS} - \nu_K$ relation, the blue dotted curve depicts the contribution from the angular momentum (we plot the series (28) up to \mathcal{V}^3), for the blue dot-dashed curve we added the first contribution from the quadrupole moment M_2 (we cut the series at \mathcal{V}^4), for the dashed blue line we added the first contribution from the octupole mass-current (series expansion up to \mathcal{V}^7) and finally in the continuous blue line we consider

⁶ Because the scale of the ν_ρ^P and ν_z^P frequencies are very similar, we decided to plot in Fig. 7 ν_ρ^P and ν_z^{OS} whose scales are different allowing a more clear comparison with the PRS solution in a single figure.

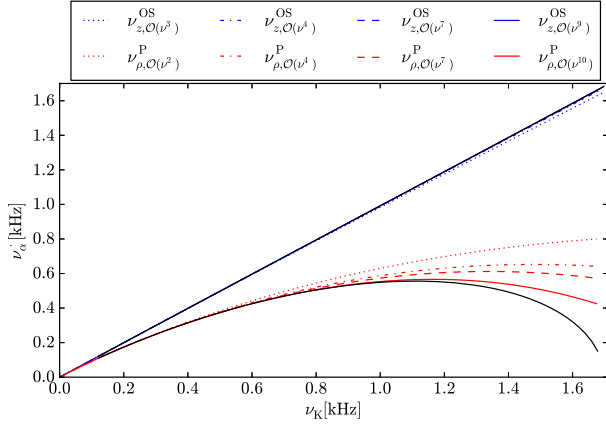


Fig. 7.— Comparison of the $\nu_z^{\text{OS}}-\nu_K$ and $\nu_\rho^{\text{P}}-\nu_K$ relations given by the PRS solution and the approximate expressions (27–28) derived by Ryan (1995). The lowest multipole moments M_0 , J , M_2 , and S_3 have been fixed to the values of the Model 2 of the Table VI of Pappas and Apostolatos (2012): $M_0 = 2.071$ km ($1.402M_\odot$), $j = 0.194$, $M_2 = -2.76$ km³ (3.73×10^{43} g cm²), and $S_3 = -2.28$ km⁴.

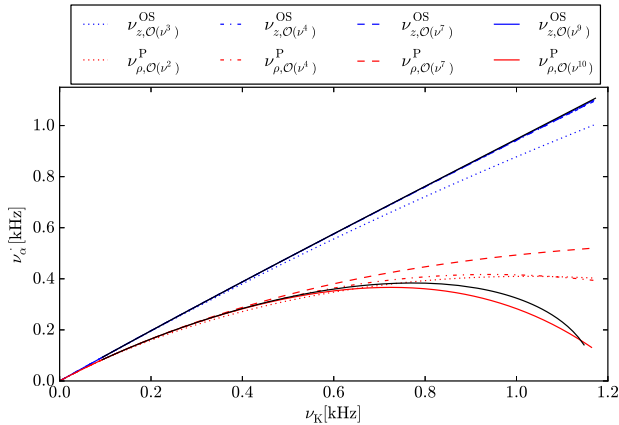


Fig. 8.— Comparison of the $\nu_z^{\text{OS}}-\nu_K$ and $\nu_\rho^{\text{P}}-\nu_K$ relations as given by the PRS solution and the approximate expressions (27–28) derived by Ryan (1995). The lowest multipole moments M_0 , J , M_2 , and S_3 have been fixed to the values of the Model 20 of the Table VI of Pappas and Apostolatos (2012): $M_0 = 4.167$ km ($2.82M_\odot$), $j = J/M_0^2 = 0.70$, $M_2 = -79.8$ km³ (-1.08×10^{45} g cm²) and $S_3 = -401.0$ km⁴.

contributions for higher multipole moments and stop the series at the order \mathcal{V}^9 , not shown in Eqs. (28). For this case, we can see that Ryan’s expressions clearly tend, from the bottom, to the exact result (continuous black curve) obtained by using the PRS solution.

For the analysis of the $\nu_\rho^{\text{P}}-\nu_K$ relation we followed the same procedure as described above. In this case, the Ryan’s expressions tend from the top to the exact result, the continuous black curve, represented by the PRS solution. It is interesting to see that the introduction of the octupole moment (dashed red line) makes the approximation to deviate from the exact result, however by including more terms the accuracy is enhanced. As can be seen from Figs. 7 and 8 the quantitative accuracy of the Ryan’s approximate formulas in the periastron precession frequency ν_ρ^{P} is less than the one obtained in the vertical oscillation frequency ν_z^{OS} .

The importance of the high-order multipole moments such as the quadrupole and the octupole moments is evident in the high-frequency regime. This is in line with the results shown

in Figs. 2–3 and in Figs. 4–5. We can see from Figs. 7 and 8 that the Ryan’s approximate formulas describe more accurately the Model 2 than the Model 20. The reason is that, as we mentioned above, we should expect a better accuracy of the series expansions from low to moderate moderate rotation rates and consequently the same occur for the quadrupole deformations. It is clear that there are appreciable differences both in rotation and deformation between the two selected models; we recall also that the rotation frequency of the star can be expressed as a function of the dimensionless j parameter as $\nu_s = G j M_0^2 / (2\pi c I) = 1.4(M/M_\odot)^2 / I_{45}$ kHz.

It is noteworthy that we have checked that the Ryan’s series expansions, Eqs. (27) and (28), fit quite accurately the exact results if taken up to order \mathcal{V}^{10} . In particular the values of the vertical oscillation and precession frequencies are fit better than the corresponding radial ones. For the Model 2 the radial oscillation frequency is well fitted by the Ryan’s expression up to Keplerian frequencies of order ~ 1.2 kHz while, for the Model 20, the approximate formulas break down at a lower value ~ 0.7 kHz. These results are of particular relevance because it makes possible the extraction of the object parameters (for instance the lowest multipoles up to S_3) by the fitting of the observed QPO frequencies in LMXBs, providing they are indeed related to the precession and oscillation frequencies of matter in the accretion disk (see Section 6, for details) and for Keplerian motion not exceeding a few kHz of frequency.

5. ACCURACY OF PRS SOLUTION

We turn now to analyze the behavior of the Kerr and PRS solutions in predicting results for the Keplerian, frame dragging, and vertical oscillation frequencies, for realistic NSs. In particular, we compare their predictions with the frequencies calculated by Morsink and Stella (1999). Since Morsink and Stella (1999) did not include the values of the octupole current moment S_3 , here we set $s = 0$ in Eq. (10) for the PRS solution. For the sake of comparison, we choose the results derived by Morsink and Stella (1999) for the EoS L, because for this EoS the highest rotating parameter j and quadrupole moment M_2 were found. In addition, the stiffness of such an EoS allows the maximum mass of the NS to be larger than the highest observed NS mass, $M_0 = 1.97 \pm 0.04M_\odot$, corresponding to the 317 Hz (3.15 milliseconds rotation period) pulsar J1614-2230 (see Demorest et al. 2010, for details).

This regime of high j and M_2 in realistic models is particularly interesting to test the deviations of the Kerr solution in the description of NS signatures as well as to explore the accuracy of the PRS solution. In Table 2, we present the results for four different sets of the star spin frequency ν_s , namely $\nu_s = 290$ Hz (M1 and M2), $\nu_s = 360$ Hz (M3 and M4), $\nu_s = 580$ Hz (M5 and M6) and $\nu_s = 720$ Hz (M7 and M8).

In Table 2, we clearly observe that the results predicted by the PRS_{s=0} solution for the Keplerian and frame-dragging frequencies are in excellent agreement with those calculated by Morsink and Stella (1999) for even highly massive, rotating and deformed models such as the model M7 with $M_0 = 2.17M_\odot$, $j = 0.51$ and $M_2 = -39.4Q_0$. We notice that Morsink and Stella (1999) reported some configurations with negative values of ν_z (see Table 2). We advance the possibility that this is due to instabilities of the numerical code that occur when the ISCO radius is located very close or inside the surface of the object. Thus, the values of the frequencies given by the analytic solution in these cases are to be considered predic-

Model	M_0/M_\odot	j	M_2/Q_0	r_+^{MS} [km]	r_+^{Kerr}	r_+^{PRS}	$\nu_{\text{K}}^{\text{MS}}$ [kHz]	$\nu_{\text{K}}^{\text{Kerr}}$	$\nu_{\text{K}}^{\text{PRS}}$	$\nu_{\text{LT}}^{\text{MS}}$ [Hz]	$\nu_{\text{LT}}^{\text{Kerr}}$	$\nu_{\text{LT}}^{\text{PRS}}$	$\nu_z^{\text{P,MS}}$ [Hz]	$\nu_z^{\text{P,Kerr}}$	$\nu_z^{\text{P,PRS}}$
M1	1.88	0.19	-5.3	15.4	14.90	15.42	1.31	1.363	1.304	39.7	42.248	38.476	18.6	39.697	22.040
M2	2.71	0.14	-3.0	22.2	22.16	22.23	0.90	0.906	0.902	19.6	19.676	19.493	17.2	18.809	17.629
M3	1.94	0.24	-8.2	15.6	14.89	15.63	1.29	1.380	1.296	49.8	57.001	49.804	17.5	52.608	26.197
M4	2.71	0.18	-4.8	21.8	21.62	21.74	0.93	0.937	0.931	26.1	27.245	26.833	21.9	25.670	23.635
M5	2.07	0.40	-23.1	16.3	14.18	16.06	1.26	1.514	1.289	84.3	125.75	88.905	-10.5	109.04	31.140
M6	2.72	0.30	-13.9	20.6	20.05	20.45	1.01	1.041	1.015	53.5	57.467	54.391	35.8	51.861	42.854
M7	2.17	0.51	-39.4	17.0	13.58	16.53	1.22	1.637	1.269	106.7	201.52	116.25	-51.8	166.55	29.804
M8	2.73	0.39	-24.4	19.8	18.85	19.65	1.06	1.136	1.077	78.8	90.821	80.895	38.4	79.079	57.187

TABLE 2

ISCO RADIUS r_+ , KEPLERIAN FREQUENCY ν_{K} , FRAME-DRAGGING (LENSE-THIRING) FREQUENCY ν_{LT} , AND VERTICAL PRECESSION FREQUENCY ν_z^{P} OF THE CO-ROTATING ORBITS CALCULATED NUMERICALLY BY MORSINK AND STELLA (1999) (UPPER INDEX MS) AND COMPARISON WITH THE CORRESPONDING PREDICTED VALUES GIVEN BY THE KERR (UPPER INDEX KERR) AND THE PRS_{S=0} SOLUTION (UPPER INDEX PRS). THE QUADRUPOLE MOMENT M_2 HAVE BEEN NORMALIZED FOR CONVENIENCE TO THE VALUE $Q_0 = 10^{43} \text{ g cm}^2$.

tions to be tested for future numerical computations. This fact can be checked within the calculations of Morsink and Stella (1999) by exploring the properties of counter-rotating orbits which produce in general ISCO radii larger than the ones of the corotating ones. In Table 3, we depicted the results in the counter-rotating case where we can notice an improvement of the accuracy of the PRS solution with respect to the corotating case.

In this line, we consider worth performing numerical computations of the precession and oscillation frequencies of particles around realistic NSs in a wider space of parameters and using up-to-date numerical techniques which will certainly help to establish and elucidate more clearly the accuracy of analytic models. It is also appropriate recalling the recent results of Pappas and Apostolatos (2012) on the computation of the general relativistic multipole moments in axially symmetric spacetimes.

6. THE RELATIVISTIC PRECISION MODEL

The X-ray light curves of LMXBs show a variability from which a wide variety of QPOs have been measured, expanding from relatively low \sim Hz frequencies all the way up to high \sim kHz frequencies (see e.g. van der Klis 1995, for details). In particular, such frequencies usually come in pairs (often called twin peaks), the lower and upper frequencies, ν_l and ν_h respectively. BHs and NSs with similar masses can show similar signatures and therefore the identification of the compact object in a LMXB is not a simple task. If the QPO phenomena observed in these systems are indeed due to relativistic motion of accretion disk matter, the knowledge of the specific behavior of the particle frequencies (e.g. rotation, oscillation, precession) in the exterior geometry of NSs and BHs becomes essential as a tool for the identification of the nature of the compact object harbored by a LMXB.

It is not the scope of this work to test a particular model for the QPO phenomenon in LMXBs but instead to show the influence of the high multipole moments on the orbital motion of test particles especially the role of the quadrupole moment which is of particular interest to differentiate a NS from a BH. There are in the literature several models that describe the QPOs in LMXBs through the frequencies of particles around the compact object, and for a recent review and comparison of the different models we refer to the recent work of Lin et al. (2011). In order to show here the main features and differences between the Kerr and the PRS solutions we shall use the Relativistic Precession Model (RPM).

The RPM model identifies the lower and higher (often called twin-peaks) kHz QPO frequencies, ν_l and ν_h , with the periastron precession and Keplerian frequencies, namely $\nu_l = \nu_\rho^{\text{P}}$ and $\nu_h = \nu_{\text{K}}$, respectively. The so-called horizontal branch oscillations (HBOs), which belong to the low frequency QPOs observed in high luminosity Z-sources (see e.g.

van der Klis 1995, for details), are related within the RPM model to the nodal precession frequency ν_z^{P} of the same orbits (Morsink and Stella 1999, see). We will use here in particular the realistic NS models of Morsink and Stella (1999) for the EoS L.

One of the salient features of the RPM model is that in the case of the HBO frequencies, the relations inferred from the first term of the expansions (27) and (28)

$$\nu_{\text{K}} = 3^{-3/5} (2\pi)^{-2/5} m^{-2/5} (\nu_\rho^{\text{P}})^{3/5}, \quad (29)$$

$$\nu_z^{\text{P}} = (2/3)^{6/5} \pi^{1/5} j m^{1/5} (\nu_\rho^{\text{P}})^{6/5}, \quad (30)$$

which implies a nodal precession frequency proportional to the square of the Keplerian frequency has been observed in some sources, for instance in the LMXB 4U 1728–34 (see Ford and van der Klis 1998, for details). In addition, 6/5 power law relating the nodal and periastron precession frequencies can explain (see Stella et al. 1999) the correlation between two of the observed QPO frequencies found in the fluxes of NSs and BHs LMXBs (see Psaltis et al. 1999, for details). This fact provides, at the same time, a significant test of the Ryan’s analytic expressions.

It is interesting to analyze the level of predictability of the precession and oscillation frequencies on particular astrophysical sources. In Fig. 9 we show the ν_l – ν_h relation within the RPM model, namely ν_ρ^{P} versus ν_{K} for the models M1–M8 of Table 2. In the upper panel we show the results for the PRS solution while, in the lower panel, we present the results for the Kerr solution. We have indicated the QPO frequencies observed in the sources GX 5–1 (see e.g. Wijnands et al. 1998; Jonker et al. 2002), 4U 1735–44 (see e.g. Ford et al. 1998), 4U 1636–53 (see e.g. Wijnands et al. 1997), Sco X1 (see e.g. van der Klis et al. 1996), GX 17–2 (see e.g. Homan et al. 2002), GX 340+0 (see e.g. Jonker et al. 2000), Cir X1 (see e.g. van der Klis et al. 1996), 4U 0614+091 (see e.g. Ford et al. 1997), and 4U 1728–34 (see e.g. Strohmayer et al. 1996).

Both the upper and lower panels of Fig. 9 have been plotted using the same frequency scales in order to aid the identification of the differences between the Kerr and the PRS solutions. One can notice that all the solid curves in the Kerr solution (lower panel of Fig. 9) are outside the range of the observed QPO frequencies exemplified, while all dashed and solid curves of the PRS are inside the QPO range. It is then clear that making a fit of the observed QPO frequencies of the selected LMXBs of Fig. 9 will necessarily require a different choice of parameters in the Kerr and PRS solutions. Therefore, conclusions for instance on the NS parameters (e.g. mass, angular momentum, quadrupole deformation) based on fitting QPOs using the Kerr geometry will deviate from the actual parameters (see e.g. Laarakkers and Poisson 1999, for details), extractable more reliably from a more complex geometry, such as the PRS one, that allows a better es-

Model	$r_-^{\text{MS}}[\text{km}]$	r_-^{Kerr}	r_-^{PRS}	$\nu_K^{\text{MS}}[\text{kHz}]$	ν_K^{Kerr}	ν_K^{PRS}	$\nu_{\text{IT}}^{\text{MS}}[\text{Hz}]$	$\nu_{\text{IT}}^{\text{Kerr}}$	$\nu_{\text{IT}}^{\text{PRS}}$	$\nu_z^{\text{MS}}[\text{Hz}]$	ν_z^{Kerr}	ν_z^{PRS}
M1	18.8	18.35	18.73	0.99	1.023	0.997	21.7	22.61	21.39	29.0	23.88	29.70
M2	25.9	25.82	25.88	0.73	0.734	0.732	12.4	12.43	12.36	13.6	12.95	13.46
M3	19.9	19.39	19.89	0.93	0.960	0.928	24.0	25.79	24.04	32.9	27.61	34.20
M4	26.4	26.33	26.42	0.71	0.715	0.712	14.7	15.08	14.94	16.4	15.88	16.57
M5	23.3	22.22	23.24	0.77	0.816	0.768	28.2	32.59	28.78	41.3	36.31	44.65
M6	28.2	27.94	28.18	0.65	0.660	0.652	20.5	21.19	20.70	24.1	23.03	24.51
M7	25.9	24.37	25.78	0.67	0.731	0.678	28.9	34.62	29.64	43.4	39.56	47.90
M8	29.7	29.18	29.58	0.61	0.623	0.611	23.2	24.40	23.48	28.4	27.12	29.22

TABLE 3
SAME AS IN TABLE 2, BUT FOR THE COUNTER-ROTATING CASE.

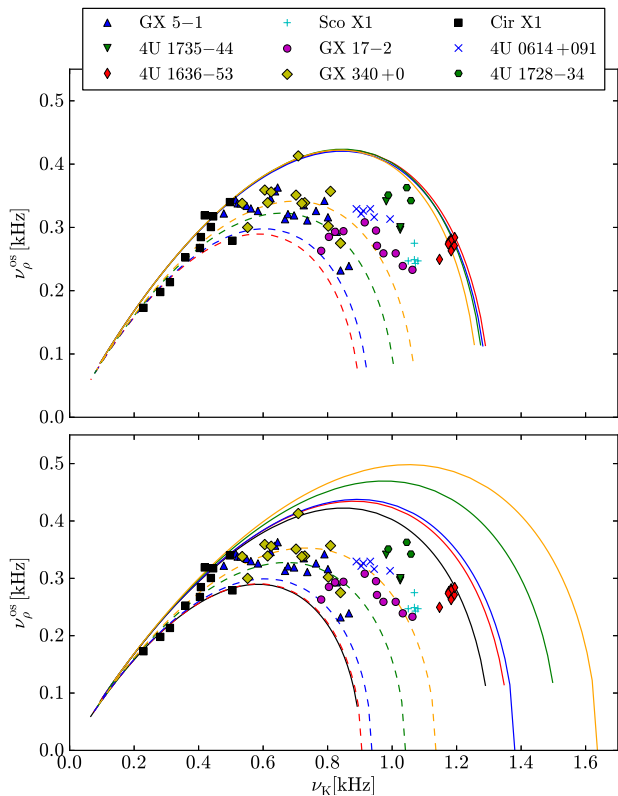


FIG. 9.— Periastron oscillation frequency, ν_ρ^{OS} , as a function of the Keplerian frequency ν_K for the NS realistic models in Table 2. We indicate the QPO frequencies observed in the sources GX 5–1, 4U 1735–44, 4U 1636–53, Sco X1, GX 17–2, GX 340+0, Cir X1, 4U 0614+091, and 4U 1728–34. The solid curves depict the results for the models M1 (solid) and M2 (dashed) with red lines, for the models M3 (solid) and M4 (dashed) with blue lines, for the models M5 (solid) and M6 (dashed) with green lines while orange lines stands for the results from models M7 (solid) and M8 (dashed). In the upper panel we present the results derived from the PRS_{s=0} solution while in the lower panel we present the results for the Kerr solution. In the lower panel we have added, to guide the eye, the inner red dashed and outer red solid curves of the upper panel using black lines.

timate for instance of the quadrupole moment of a compact star.

In Fig. 9 we show the relation ν_z^{P} versus ν_K for the models M1–M8 of Table 2. For the sake of comparison we show the low frequency branch observed in the LMXB 4U 1728–34 (see Ford and van der Klis 1998, for details). From the analysis of the pulsating X-ray flux it turns out that very likely the spin frequency of the NS in 4U 1728–34 is ~ 363 Hz (see Strohmayer et al. 1996, for details). Thus, the models M3 ($M_0 = 1.94M_\odot$, $j = 0.24$) and M4 ($M_0 = 2.71M_\odot$, $j = 0.18$) in Table 2 that correspond to a NS of spin frequency 360 Hz are of particular interest for the analysis of this source. It was suggested by Stella et al. (1999); Stella and Vietri (1999) that the low frequency observed in 4U 1728–34 are likely to be due

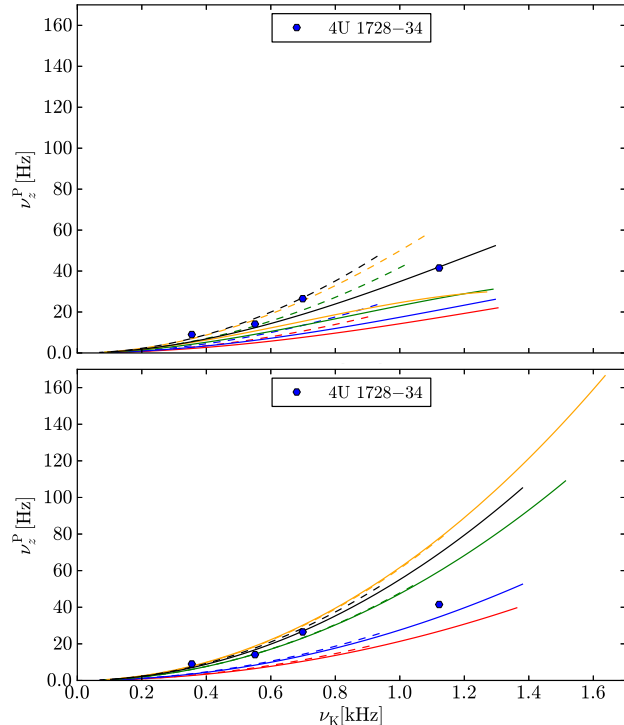


FIG. 10.— Nodal precession frequency, ν_z^{P} , as a function of the Keplerian frequency ν_K for the NS realistic models in Table 2. The convention is as Fig. 9. We indicate the QPO frequencies observed in the LMXB 4U 1728–34 (see Ford and van der Klis 1998). The black curves indicate the $2\nu_z^{\text{P}} - \nu_K$ relation for the models M3 and M4 (solid and dashed) following the suggestion of Stella et al. (1999); Stella and Vietri (1999).

to excitations of the second harmonic of the vertical motion and therefore a better fit of the lower-higher QPO frequencies of 4U 1728–34 (and of similar sources) will be obtained for the relation $2\nu_z^{\text{P}} - \nu_K$. The black curves in Fig. 10 indicate the $2\nu_z^{\text{P}} - \nu_K$ relation for the models M3 and M4 (solid and dashed) following the above suggestion. Although the improvement of the fit is evident, we notice that the NS parameters that correctly reproduce the features of 4U 1728–34 are likely in between the models M3 and M4.

7. CONCLUDING REMARKS

We have done an extensive comparison of the orbital motion of neutral test particles in the PRS and Kerr spacetime geometries. In particular we have emphasized on the Keplerian and frame-dragging frequencies, as well as the precession and oscillation frequencies of the radial and vertical motions.

We have evidenced the differences in this respect between the Kerr and PRS solution, especially in the rapid \sim kHz rotation regime. Such differences are the manifestation of the influence of the high order multipole moments such as the quadrupole and octupole.

The analysis of the deviations between the Kerr and PRS features for given mass and angular momentum of a source studied in this work are useful to distinguish the signatures between BHs and NSs, which relevant to establish a separatrix for the identification of the compact objects harbored in X-Ray Binaries. In the case of BH candidates, these results might become important for testing the no-hair theorem of BHs (see e.g. [Johannsen and Psaltis 2011](#)). Equally important, the application of the precession and oscillation frequencies to the explanation of QPOs in LMXBs possessing a NS, can unveil information on the NS parameters, leading to a possible identification of the behavior of the nuclear matter EoS at supranuclear densities. In this line, the identification of the rotation frequency of NSs in LMXBs from the pulsating X-ray flux ν_{burst} , e.g. the case of 4U 1728–34 ([Ford and van der Klis 1998](#)), 4U 1916–053 ([Galloway et al. 2001](#)) and more recently the case of IGR J17191–2821 ([Altamirano et al. 2010](#)), will certainly help to constrain QPO models as well as the NS parameters. Additional information coming from the photospheric radius expansion phenomena observed in these systems (see e.g. [Muno et al. 2001](#), for details) during their transient activity with Super-Eddington emission can become of paramount importance if combined with the QPO information.

The generalization of the present work to the electrovacuum case is important to establish the influence of the magnetic dipole and quadrupole moments on the orbital motion of particles around compact objects (see e.g. [Bakala et al. 2010](#); [Sanabria-Gómez et al. 2010](#); [Bakala et al. 2012](#)).

Interesting effects on the epicyclic frequencies due to the presence of the magnetic dipole have been already pointed out recently by [Bakala et al. \(2010\)](#) and [Bakala et al. \(2012\)](#). These effects were predicted after neglecting the contribution of the electromagnetic field to the curvature, for $j = 0$ see [Bakala et al. \(2010\)](#) and for $j \neq 0$ [Bakala et al. \(2012\)](#). [Bakala et al. \(2010\)](#) assumed the model of the star as a dipole magnetic field superposed on a Schwarzschild black hole. In the second work, they studied the case of a magnetized slowly rotating neutron stars; to build the model they super-

pose an dipolar magnetic field on the Lense-Thirring geometry. The effects of the magnetic dipole on the location of the ISCO, within the PRS solution, has been investigated by [Sanabria-Gómez et al. \(2010\)](#).

A complete analysis of the effects due to the emergence of electromagnetic structure on the orbital motion of charged particles is therefore of interest and deserve the appropriate attention. Recent observations have shown that for stars with strong magnetic fields the quadrupole and octupole magnetic terms make significant contributions to the magnetic field ([Donati et al. 2006](#)), which indicates that arbitrary higher order multipole components might be required in a realistic model. The presence of a magnetic quadrupole demands the breaking of the reflection symmetry (see [Pachón and Sanabria-Gómez 2006](#), for details), by means of a slightly change to the Ernst electric potential over the symmetry axis

$$f(z) = \frac{qz^2 + i\mu z + i\zeta}{z^3 + z^2(m - ia) - kz + is}, \quad (31)$$

a quadrupolar magnetic component $\mathcal{B}_2 = \zeta$ can be introduced to the PRS solution. Such a change generates just a redefinition of the coefficients f_i in Eq. (15). In this way the PRS solution can be readily use to explore the effect of strong magnetic fields with non-dipolar structure.

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APPENDIX

METRIC FUNCTIONS

The functions A , B , C , H , G , K , and I used to express the metric functions (17) are given by

$$A = \sum_{1 \leq i < j < k \leq 6} a_{ijk} r_i r_j r_k, \quad B = \sum_{1 \leq i < j \leq 6} b_{ij} r_i r_j, \quad C = \sum_{1 \leq i < j \leq 6} c_{ij} r_i r_j, \quad K = \sum_{1 \leq i < j < k \leq 6} a_{ijk}, \quad (\text{A1})$$

$$H = zA - (\beta_1 + \beta_2 + \beta_3)B + \sum_{1 \leq i < j < k \leq 6} h_{ijk} r_i r_j r_k + \sum_{1 \leq i < j \leq 6} (\alpha_i + \alpha_j) b_{ij} r_i r_j, \quad (\text{A2})$$

$$G = -(\beta_1 + \beta_2 + \beta_3)A + zB + \sum_{1 \leq i < j \leq 6} g_{ij} r_i r_j + \sum_{1 \leq i < j < k \leq 6} (\alpha_i + \alpha_j + \alpha_k) a_{ijk} r_i r_j r_k, \quad (\text{A3})$$

$$I = (f_1 + f_2 + f_3)(A - B) + (\beta_1 + \beta_2 + \beta_3 - z)C + \sum_{1 \leq i < j < k \leq 6} p_{ijk} r_i r_j r_k + \sum_{i=1}^6 p_i r_i + \sum_{1 \leq i < j \leq 6} [p_{ij} - (\alpha_i + \alpha_j) c_{ij}] r_i r_j, \quad (\text{A4})$$

with

$$\begin{aligned} r_i &= \sqrt{\rho^2 + (z - \alpha_i)^2}, \quad a_{ijk} = (-1)^{i+j+1} \Lambda_{ijk} \Gamma_{lmn}, \quad b_{ij} = (-1)^{i+j} \lambda_{ij} H_{lmnp}, \\ c_{ij} &= (-1)^{i+j} \lambda_{ij} [f(\alpha_l) \Gamma_{mnp} - f(\alpha_m) \Gamma_{npl} + f(\alpha_n) \Gamma_{plm} - f(\alpha_p) \Gamma_{lpm}], \\ h_{ijk} &= (-1)^{i+j+k} \Lambda_{ijk} (e_1^* \delta_{23|lmn} + e_2^* \delta_{31|lmn} + e_3^* \delta_{12|lmn}), \quad g_{ij} = (-1)^{i+j} \lambda_{ij} (\alpha_l \Gamma_{mnp} - \alpha_m \Gamma_{npl} + \alpha_n \Gamma_{plm} - \alpha_p \Gamma_{lpm}), \\ p_i &= (-1)^i D_i [f(\alpha_l) H_{mnp} - f(\alpha_m) H_{npsl} + f(\alpha_n) H_{plsm} - f(\alpha_p) H_{slmn} + f(\alpha_s) H_{lmnp}], \\ p_{ij} &= (-1)^{i+j} \lambda_{ij} (e_1^* \Upsilon_{23|lmnp} + e_2^* \Upsilon_{31|lmnp} + e_3^* \Upsilon_{12|lmnp}), \quad p_{ijk} = (-1)^{i+j+1} \Lambda_{ijk} (e_1^* \Psi_{23|lmn} + e_2^* \Psi_{31|lmn} + e_3^* \Psi_{12|lmn}), \\ \lambda_{ij} &= (\alpha_i - \alpha_j) D_i D_j, \quad \Lambda_{ijk} = (\alpha_i - \alpha_j)(\alpha_i - \alpha_k)(\alpha_j - \alpha_k) D_i D_j D_k, \\ D_i &= \frac{1}{(\alpha_i - \beta_1)(\alpha_i - \beta_2)(\alpha_i - \beta_3)}, \quad \Gamma_{lmn} = H_3(\alpha_l) \Delta_{12|mn} + H_3(\alpha_m) \Delta_{12|nl} + H_3(\alpha_n) \Delta_{12|lm}, \end{aligned}$$

and

$$\begin{aligned}\Delta_{lm|np} &= H_l(\alpha_n) H_m(\alpha_p) - H_l(\alpha_p) H_m(\alpha_n), \quad H_l(\alpha_n) = \frac{2 \prod_{p \neq n} (\alpha_p - \beta_l^*)}{\prod_{k \neq l}^3 (\beta_l^* - \beta_k^*) \prod_{k=1}^3 (\beta_l^* - \beta_k)} - \sum_{k=1}^3 \frac{2f_l^* f_k}{(\beta_l^* - \beta_k)(\alpha_n - \beta_k)}, \\ \delta_{lm|nps} &= \Delta_{lm|np} + \Delta_{lm|ps} + \Delta_{lm|sn}, \quad h_{l|mp} = H_3(\alpha_l) \delta_{12|mp}, \quad H_{l|mp} = h_{l|mp} + h_{m|npl} + h_{n|plm} + h_{p|lmn}, \\ \Psi_{lm|nps} &= f(\alpha_n) \Delta_{lm|ps} + f(\alpha_p) \Delta_{lm|sn} + f(\alpha_s) \Delta_{lm|np}, \\ \Upsilon_{lm|nprs} &= f(\alpha_n) \delta_{lm|prs} - f(\alpha_p) \delta_{lm|rsn} + f(\alpha_r) \delta_{lm|snp} - f(\alpha_s) \delta_{lm|npr},\end{aligned}$$

being α 's the roots of the Sibgatullin equation [Sibgatullin \(1991\)](#); [Manko and Sibgatullin \(1993\)](#)

$$e(z) + \tilde{e}(z) + 2\tilde{f}(z)f(z) = 0. \quad (\text{A5})$$

KERR'S METRIC IN WEYL-PAPAPETROU QUASI-CYLINDRICAL COORDINATES

In order to keep comparisons in the save place, we consider useful to display the Kerr solution in the Weyl-Papapetrou quasi-cylindrical coordinates. For this case,

$$f = \frac{A\bar{A} - B\bar{B}}{(A - B)(\bar{A} - \bar{B})}, \quad e^{2\gamma} = \frac{A\bar{A} - B\bar{B}}{K\bar{K} \prod_{n=1}^2 r_n}, \quad \omega = \frac{\text{Im}[(A + B)\bar{H} - (\bar{A} + \bar{B})G]}{A\bar{A} - B\bar{B}}, \quad (\text{B1})$$

where for our own convenience we do not present the definition of each term, but present the final combination of them, i.e.,

$$\begin{aligned}A\bar{A} - B\bar{B} &= -8(a^2 - m^2)^3 (\rho^2 + z^2) \left(m^2 \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2} \right. \\ &\quad \left. - 2a^2 \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2 + a^2 m^2 - m^4 + m^2 \rho^2 + m^2 z^2} \right),\end{aligned} \quad (\text{B2})$$

$$\begin{aligned}(A - B)(\bar{A} - \bar{B}) &= -8(m^2 - a^2)^3 \\ &\quad \times (\rho^2 + z^2) \left(a^2 \left(2m \left(\sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} + \sqrt{2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \right) \right. \right. \\ &\quad \left. \left. + 2\sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \sqrt{2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} + 3m^2 \right) \right. \\ &\quad \left. - m^2 \left(2m \left(\sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} + \sqrt{2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \right) \right. \right. \\ &\quad \left. \left. + \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \sqrt{2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} + 3m^2 + \rho^2 + z^2 \right) \right),\end{aligned} \quad (\text{B3})$$

$$K\bar{K} \prod_{n=1}^2 r_n = 16(m^2 - a^2)^4 (\rho^2 + z^2) \sqrt{(z - \sqrt{m^2 - a^2})^2 + \rho^2} \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2} \quad (\text{B4})$$

$$\begin{aligned}\text{Im}[(A + B)\bar{H} - (\bar{A} + \bar{B})G] &= 16am(m^2 - a^2)^3 (\rho^2 + z^2) \left(-m^2 \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \right. \\ &\quad \left. - m \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2} \right. \\ &\quad \left. + a^2 \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} - z \sqrt{m^2 - a^2} \sqrt{-2z\sqrt{m^2 - a^2} - a^2 + m^2 + \rho^2 + z^2} \right. \\ &\quad \left. - m^2 \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2} + a^2 \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2} \right. \\ &\quad \left. + z \sqrt{m^2 - a^2} \sqrt{(\sqrt{m^2 - a^2} + z)^2 + \rho^2} + a^2 m - m^3 + m\rho^2 + m^2 z^2 \right).\end{aligned} \quad (\text{B5})$$

From here, it is clear how changing $a \rightarrow -a$ will cause only a global change in the sign of the metric function ω and therefore only a change in the $g_{t\phi}$ metric component.